

11-5-2018

## The Shape of the Orbit in FLRW Spacetimes

David Garfinkle

*Oakland University*, garfinkl@oakland.edu

Lawrence R. Mead

*University of Southern Mississippi*, Lawrence.Mead@usm.edu

H.I. Ringermacher

*University of Southern Mississippi*, ringerha@gmail.com

Follow this and additional works at: [https://aquila.usm.edu/fac\\_pubs](https://aquila.usm.edu/fac_pubs)



Part of the [Astrophysics and Astronomy Commons](#)

---

### Recommended Citation

Garfinkle, D., Mead, L. R., Ringermacher, H. (2018). The Shape of the Orbit in FLRW Spacetimes. *Journal of Physics Communications*, 2(11), 1-5.

Available at: [https://aquila.usm.edu/fac\\_pubs/15777](https://aquila.usm.edu/fac_pubs/15777)

This Article is brought to you for free and open access by The Aquila Digital Community. It has been accepted for inclusion in Faculty Publications by an authorized administrator of The Aquila Digital Community. For more information, please contact [Joshua.Cromwell@usm.edu](mailto:Joshua.Cromwell@usm.edu).

LETTER • OPEN ACCESS

## The shape of the orbit in FLRW spacetimes

To cite this article: David Garfinkle *et al* 2018 *J. Phys. Commun.* **2** 111001

View the [article online](#) for updates and enhancements.



## LETTER

## The shape of the orbit in FLRW spacetimes

## OPEN ACCESS

RECEIVED  
24 September 2018REVISED  
9 October 2018ACCEPTED FOR PUBLICATION  
24 October 2018PUBLISHED  
5 November 2018David Garfinkle<sup>1,2,4</sup> , Lawrence R Mead<sup>3</sup> and H I Ringerha<sup>3</sup><sup>1</sup> Department of Physics, Oakland University, Rochester, MI 48309, United States of America<sup>2</sup> Michigan Center for Theoretical Physics, Randall Laboratory of Physics, University of Michigan, Ann Arbor, MI 48109-1120, United States of America<sup>3</sup> Department of Physics and Astronomy, University of Southern Mississippi, Hattiesburg, MS 39406, United States of America<sup>4</sup> Author to whom any correspondence should be addressed.E-mail: [garfinkl@oakland.edu](mailto:garfinkl@oakland.edu), [lawrence.mead@usm.edu](mailto:lawrence.mead@usm.edu) and [ringerha@gmail.com](mailto:ringerha@gmail.com)

Keywords: geodesic motion, cosmological spacetimes, symmetries, coordinates

Original content from this work may be used under the terms of the [Creative Commons Attribution 3.0 licence](https://creativecommons.org/licenses/by/4.0/).

Any further distribution of this work must maintain attribution to the author(s) and the title of the work, journal citation and DOI.



## Abstract

The shape of the orbit of a free particle is examined in Friedmann-Lemaître-Robertson-Walker (FLRW) cosmologies. The spatial projection of the orbit is time-independent and has a simple geometric description. We relate this description to the expression in terms of standard FLRW coordinates.

The spacetime of Friedmann-Lemaître-Robertson-Walker (FLRW) cosmology consists of a constant curvature space that evolves in time through a time dependent scale factor. An object in free fall, or a light ray, travels along a spacetime geodesic, which traces out a path in the constant curvature 3-space. It is a remarkable fact that this path is also a geodesic of the underlying constant curvature 3-space, and is therefore completely independent of the behavior of the scale factor. Due to the symmetries of FLRW, there are a variety of geometric and coordinate based techniques that can be used to study geodesics. The standard textbook approach[1–7] actually uses a combination of techniques: standard FLRW coordinates are used, but one makes use of homogeneity to demand that the geodesic passes through the point  $r = 0$  and is thus a radial geodesic. This is convenient because the geodesics we are most interested in are those that reach our position, and we adopt for ourselves the coordinate privileged position  $r = 0$ . Nonetheless, it is sometimes of interest to consider non-radial geodesics in FLRW, for example in the Sunyaev-Zeldovich effect[8] where cosmic background photons inverse Compton scatter off the hot gas in a galaxy cluster; or more generally when treating the cosmological Boltzmann equation (see e.g. [9]). In this note we will look at FLRW geodesics in general (i.e. not necessarily radial) form, using two different approaches: (1) a geometric approach and (2) a coordinate based approach. We will present and compare these two approaches.

The metric  $g_{ab}$  of the FLRW cosmology can be written as

$$g_{ab} = -u_a u_b + a^2 \bar{h}_{ab} \quad (1)$$

Here  $u^a = (\partial/\partial t)^a$  is the four-velocity of the FLRW fluid,  $a(t)$  is the scale factor, and  $\bar{h}_{ab}$  is the metric of a unit constant curvature space. Note that this constant curvature space can be flat, positively curved (3-sphere), or negatively curved (hyperbolic space). Let  $k^a$  be the four-velocity of a timelike or null geodesic (i.e. an object in free-fall, or a light ray). Then  $k^a k_a = -\kappa$  where  $\kappa = 1$  for material objects and  $\kappa = 0$  for light rays.  $k^a$  takes the form

$$k^a = \alpha u^a + \beta v^a \quad (2)$$

where  $v^a$  is a unit vector in the metric  $\bar{h}_{ab}$ . Since  $u^a$  is a unit timelike vector orthogonal to  $v^a$ , we have  $\alpha = -k^a u_a$ . From the geodesic equation and the standard result  $\nabla_a u_b = a\dot{a}\bar{h}_{ab} = (\dot{a}/a)(g_{ab} + u_a u_b)$  we find

$$\begin{aligned}
 k^a \nabla_a \alpha &= k^a \nabla_a (-k^b u_b) \\
 &= -k^a k^b \nabla_a u_b \\
 &= -k^a k^b \frac{\dot{a}}{a} (g_{ab} + u_a u_b) \\
 &= \frac{\dot{a}}{a} (\kappa - \alpha^2) \\
 &= \frac{\kappa - \alpha^2}{a\alpha} k^a \nabla_a a,
 \end{aligned} \tag{3}$$

from which it follows that

$$k^a \nabla_a (\alpha^2 - \kappa) = -\frac{2}{a} (\alpha^2 - \kappa) k^a \nabla_a a. \tag{4}$$

It then follows that

$$\alpha = \sqrt{\kappa + c_0^2/a^2}, \quad \beta = c_0/a^2 \tag{5}$$

for some constant  $c_0$ . (As shown in [7] equation (5) can also be derived by first showing that  $a^4 \bar{h}_{ab}$  is a Killing tensor). Now let  $\xi^a$  be a Killing vector of  $\bar{h}_{ab}$ . Then it follows that  $\xi^a$  is also a Killing vector of the FLRW spacetime. Thus there is a constant  $c_1$  such that

$$c_1 = g_{ab} k^a \xi^b = (u_a u_b + a^2 \bar{h}_{ab})(\alpha u^a + \beta v^a) \xi^b = c_0 \bar{h}_{ab} v^a \xi^b \tag{6}$$

Thus  $v^a$ , the tangent vector to the spatial projection of the orbit, is a unit vector in the underlying unit constant curvature space whose inner product with every Killing vector of that space is a constant. It then follows that the orbit is a geodesic of that space.

We now describe the orbits for each possible curvature of space: for flat space, the orbit is a straight line. For positive curvature, the orbit is a great circle of the 3-sphere. That is, realizing the three sphere as the surface  $w^2 + x^2 + y^2 + z^2 = 1$  in a flat 4-dimensional Euclidean space, the great circle is the intersection of this surface with a plane through the origin. The analogous result holds for hyperbolic space. Realizing this space as the surface (unit hyperboloid)  $-t^2 + x^2 + y^2 + z^2 = -1$  in Minkowski spacetime, the orbit is the intersection of this surface with a plane through the origin.

We now consider the coordinate description of the orbit. It is both convenient and usual in treating FLRW spacetimes to use a single expression, subsuming all three possibilities for the curvature, to describe the spacetime. The line element is given by the expression

$$ds^2 = -dt^2 + a^2(t) \left[ \frac{dr^2}{1 - kr^2} + r^2 (d\theta^2 + \sin^2 \theta d\varphi^2) \right], \tag{7}$$

where  $k = 0$  (flat space),  $k = 1$  (3-sphere), or  $k = -1$  (hyperbolic space). The geodesic equation is

$$\frac{d^2 x^\mu}{d\lambda^2} + \Gamma_{\alpha\beta}^\mu \frac{dx^\alpha}{d\lambda} \frac{dx^\beta}{d\lambda} = 0, \tag{8}$$

where  $\lambda$  is an affine parameter. Without loss of generality, we specialize to orbits in the  $\theta = \pi/2$  plane. The geodesic equation then yields

$$\frac{d^2 t}{d\lambda^2} + a\dot{a} \left[ \frac{1}{1 - kr^2} \left( \frac{dr}{d\lambda} \right)^2 + r^2 \left( \frac{d\varphi}{d\lambda} \right)^2 \right] = 0, \tag{9}$$

$$\frac{d^2 r}{d\lambda^2} + \frac{kr}{1 - kr^2} \left( \frac{dr}{d\lambda} \right)^2 - r(1 - kr^2) \left( \frac{d\varphi}{d\lambda} \right)^2 + 2\frac{\dot{a}}{a} \frac{dr}{d\lambda} \frac{dt}{d\lambda} = 0, \tag{10}$$

$$\frac{d^2 \varphi}{d\lambda^2} + 2\frac{\dot{a}}{a} \frac{d\varphi}{d\lambda} \frac{dt}{d\lambda} + \frac{2}{r} \frac{d\varphi}{d\lambda} \frac{dr}{d\lambda} = 0. \tag{11}$$

Defining  $d/d\ell = a^2 d/d\lambda$  along the geodesics, one obtains from equations (10)–(11)

$$\frac{d^2 r}{d\ell^2} + \frac{kr}{1 - kr^2} \left( \frac{dr}{d\ell} \right)^2 - r(1 - kr^2) \left( \frac{d\varphi}{d\ell} \right)^2 = 0, \tag{12}$$

$$\frac{d^2 \varphi}{d\ell^2} + \frac{2}{r} \frac{d\varphi}{d\ell} \frac{dr}{d\ell} = 0, \tag{13}$$

thus demonstrating that the shape of the orbit is independent of the scale factor. And, in fact, equations (12)–(13) are precisely the geodesic equations for the metric  $\bar{h}_{ab}$  with  $\ell$  as the length in that metric. One solution of equation (13) is  $d\varphi/d\ell = 0$  which corresponds to radial geodesics. In this case, the first integral of equation (12) leads to  $d\ell = dr/\sqrt{1 - kr^2}$ . This standard result, which can also be read off directly from equation (7), in turn leads to the standard result that

$$r = \ell \ (k = 0), \ r = \sin \ell \ (k = 1), \ r = \sinh \ell \ (k = -1). \tag{14}$$

If  $d\varphi/d\ell \neq 0$  (non-radial geodesics) then using equation (13) and the chain rule in equation (12) we obtain

$$\frac{d^2r}{d\varphi^2} - \frac{2 - 3kr^2}{r(1 - kr^2)} \left(\frac{dr}{d\varphi}\right)^2 - r(1 - kr^2) = 0. \tag{15}$$

Defining  $u \equiv r^{-2} - k$ , equation (15) becomes

$$\frac{d^2u}{d\varphi^2} - \frac{1}{2u} \left(\frac{du}{d\varphi}\right)^2 + 2u = 0, \tag{16}$$

from which it follows that

$$\frac{d}{d\varphi} \left[ u^{-1} \left(\frac{du}{d\varphi}\right)^2 + 4u \right] = 0. \tag{17}$$

Therefore there is an integration constant  $c_2$  for which

$$u^{-1} \left(\frac{du}{d\varphi}\right)^2 + 4u = c_2. \tag{18}$$

Note that since  $r \leq 1$  for  $k = 1$ , it follows that  $u \geq 0$  for any value of  $k$ , and thus using equation (18) it follows that  $c_2 \geq 0$ . Now defining  $w \equiv u - c_2/8$  we find that equation (18) is equivalent to

$$\left(\frac{dw}{d\varphi}\right)^2 + 4w^2 = \frac{c_2^2}{16}. \tag{19}$$

Now differentiating equation (19) with respect to  $\varphi$  we obtain

$$\frac{d^2w}{d\varphi^2} + 4w = 0. \tag{20}$$

The general solution to equation (20) is

$$w = A \cos(2(\varphi + \beta)) \tag{21}$$

where  $A$  is a nonnegative constant, and  $\beta$  is a constant. Now using equation (21) in equation (19) we find that the integration constant is given by  $c_2 = 8A$  (where we have also used the fact that  $c_2$  is nonnegative). It then follows from the definition of  $w$  that

$$u = 2A \cos^2(\varphi + \beta), \tag{22}$$

And thus from the definition of  $u$  that

$$r = \frac{1}{\sqrt{k + 2A \cos^2(\varphi + \beta)}}. \tag{23}$$

We now compare equation (23) to the results of the geometric approach for the three cases:  $k = 0, 1, -1$ . For  $k = 0$  the space is flat and  $(r, \varphi)$  are plane polar coordinates. However, in this case equation (23) is easily seen to be the equation for a straight line in plane polar coordinates, so the geometric result agrees with the coordinate result.

For  $k = 1$ , the geometric result is that the orbit is a circle. However, equation (23) does not look like the equation for a circle. In fact, if  $(r, \varphi)$  were plane polar coordinates, then equation (23) would be the equation of an ellipse (see figure 1). The coordinates  $(w, x, y, z)$  for the embedding of the 3-sphere in 4 dimensional flat space are related to the FLRW coordinates  $(r, \varphi)$  by

$$(w, x, y, z) = (\sqrt{1 - r^2}, r \cos \varphi, r \sin \varphi, 0) \tag{24}$$

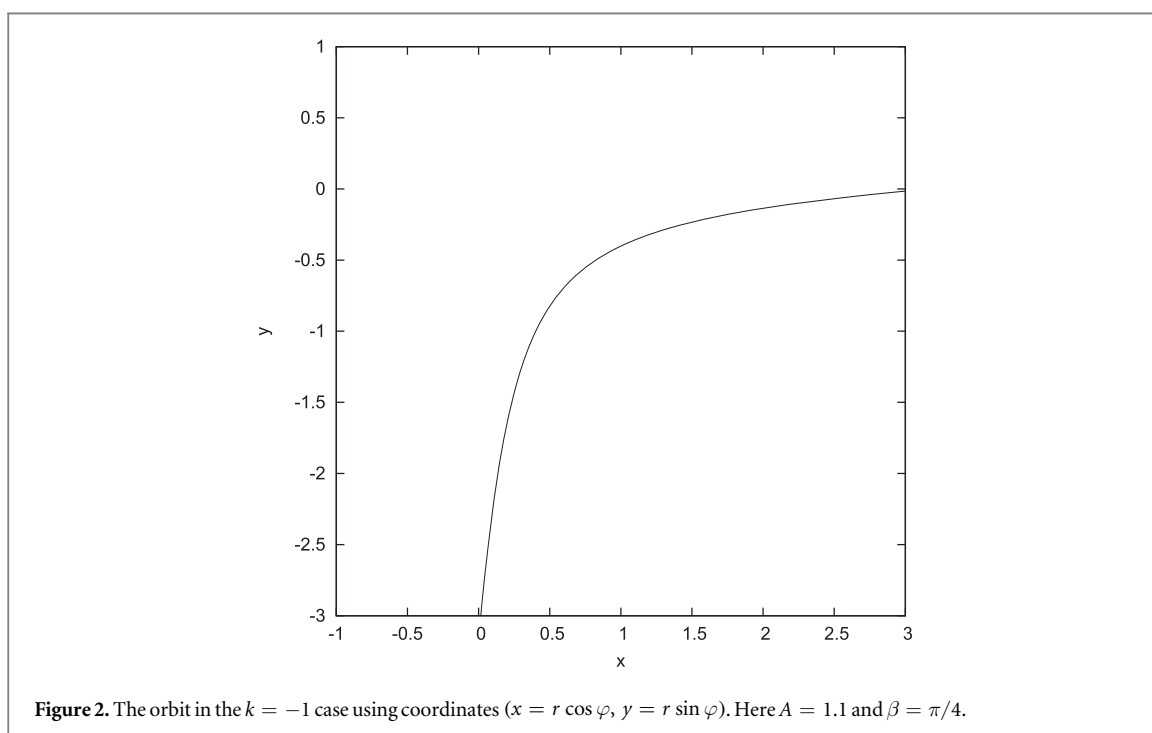
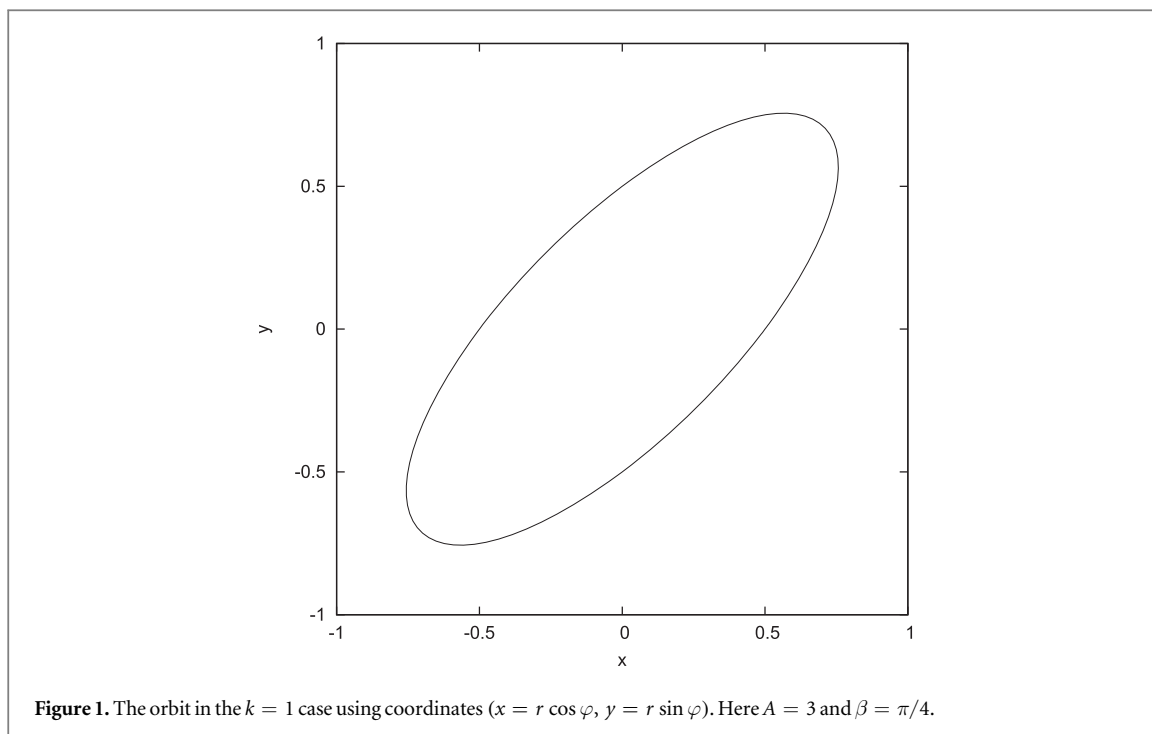
where we have also used the condition  $\theta = \pi/2$ . The equation for a plane through the origin is  $w = c_3 x + c_4 y$  for some constants  $c_3$  and  $c_4$ , which becomes using equation (24)

$$\sqrt{1 - r^2} = Br \cos(\varphi + \beta) \tag{25}$$

for some constants  $B$  and  $\beta$ . Solving for  $r$  we obtain

$$r = \frac{1}{\sqrt{1 + B^2 \cos^2(\varphi + \beta)}}. \tag{26}$$

which agrees with equation (23).



Corresponding results hold for  $k = -1$  (see figure 2). We have

$$(t, x, y, z) = (\sqrt{1 + r^2}, r \cos \varphi, r \sin \varphi, 0) \tag{27}$$

where we have also used the condition  $\theta = \pi/2$ . The equation for a plane through the origin is  $t = c_5 x + c_6 y$  for some constants  $c_5$  and  $c_6$ , which becomes using equation (27)

$$\sqrt{1 + r^2} = Br \cos(\varphi + \beta) \tag{28}$$

for some constants  $B$  and  $\beta$ . Solving for  $r$  we obtain

$$r = \frac{1}{\sqrt{-1 + B^2 \cos^2(\varphi + \beta)}}. \tag{29}$$

which agrees with equation (23).

Thus we see that for a given value of  $k$  all the FLRW orbits are the same geometrically, and are therefore equivalent to orbits of the standard radial geodesics of equation (14). However, these orbits *look* very different for non-radial geodesics expressed in standard FLRW coordinates. This is completely analogous to the well known phenomenon (familiar to airplane travelers) that though all great circles of the Earth are geometrically equivalent to the equator, they look very different when expressed in terms of longitude and latitude.

## Acknowledgments

This work was supported by NSF grants PHY-1505565 and PHY-1806219 to Oakland University.

## ORCID iDs

David Garfinkle  <https://orcid.org/0000-0003-4415-774X>

## References

- [1] Weinberg S 1972 *Gravitation and Cosmology* (New York: Wiley)
- [2] Misner C W, Thorne K S and Wheeler J A 1973 *Gravitation* (New York: W.H. Freeman and Company)
- [3] Schutz B and First Course A 2009 *General Relativity* 2nd edn (Cambridge: Cambridge University Press)
- [4] Hartle J B 2003 *Gravity* (San Francisco, CA: Addison Wesley)
- [5] Wald R M 1984 *General Relativity* (Chicago, IL: University of Chicago Press)
- [6] Ryden B 2017 *Introduction to Cosmology* 2nd edn (Cambridge: Cambridge University Press)
- [7] Carroll S 2004 *Spacetime and Geometry* (San Francisco, CA: Addison Wesley)
- [8] Sunyaev R A and Zeldovich Ya B 1980 *Mon. Not. Roy. Astron. Soc.* **190** 413
- [9] Dodelson S 2003 *Modern Cosmology* (Amsterdam: Academic)